

Max-Planck-Institut für Radioastronomie

ATLASGAL: 3-mm methanol masers in high-mass star formation regions

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Methanol (CH₃OH) masers are common phenomena in star formation regions and powerful tools to study the physical conditions of the gas. The main goals of this study are (1) to search for new methanol masers at 3 mm, (2) to statistically study the relationship between class I masers and shock tracers, (3) to study properties between methanol masers and their host clumps, also as a function of their evolutionary stage and, (4) to better constrain the physical conditions through multiple co-spatial maser line pairs.

Observations

- Targets: 408 ATLASGAL clumps
- Telescope: IRAM-30 m (83.8 – 115.7 GHz)
- Observing dates: 2010.05 – 2012.10
- Beam size: 29 arcsec @ 84 GHz
- 1 σ rms: 0.2 Jy @ ~0.7 km/s

Maser Detection

- Class I CH₃OH masers(*):
54 (**50 new**) masers at 84 GHz
100 (**29 new**) masers at 95 GHz
4 (**4 new**) masers at 104.3 GHz; the known number has increased from 5 to 9
- Class II CH₃OH masers(†):
11 (**8 new**) masers at 107 GHz; the known number has increased from 25 to 33
No sources show maser emission at 85.5, 86.6, 86.9, 104.1 and 108 GHz

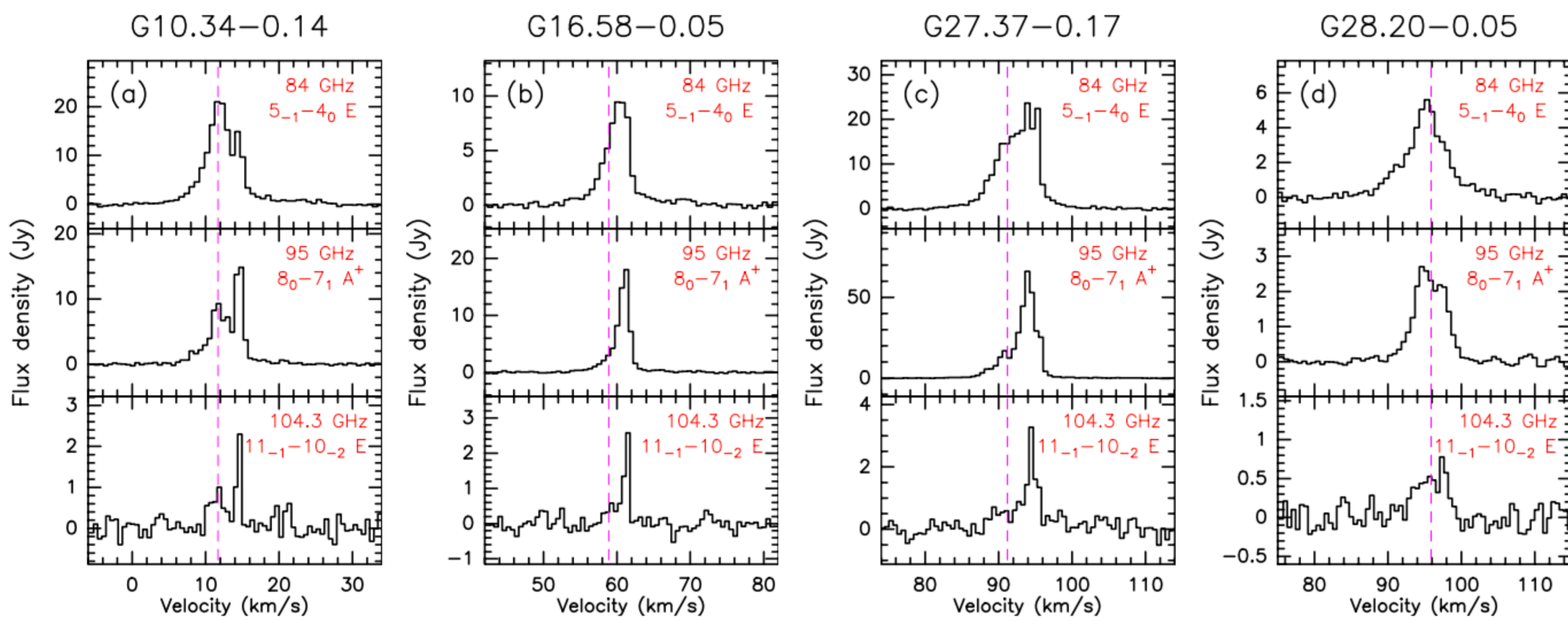


Fig. 1: Spectra of the 84, 95, 104.3 GHz lines for 4 sources showing 104.3 GHz methanol maser features.

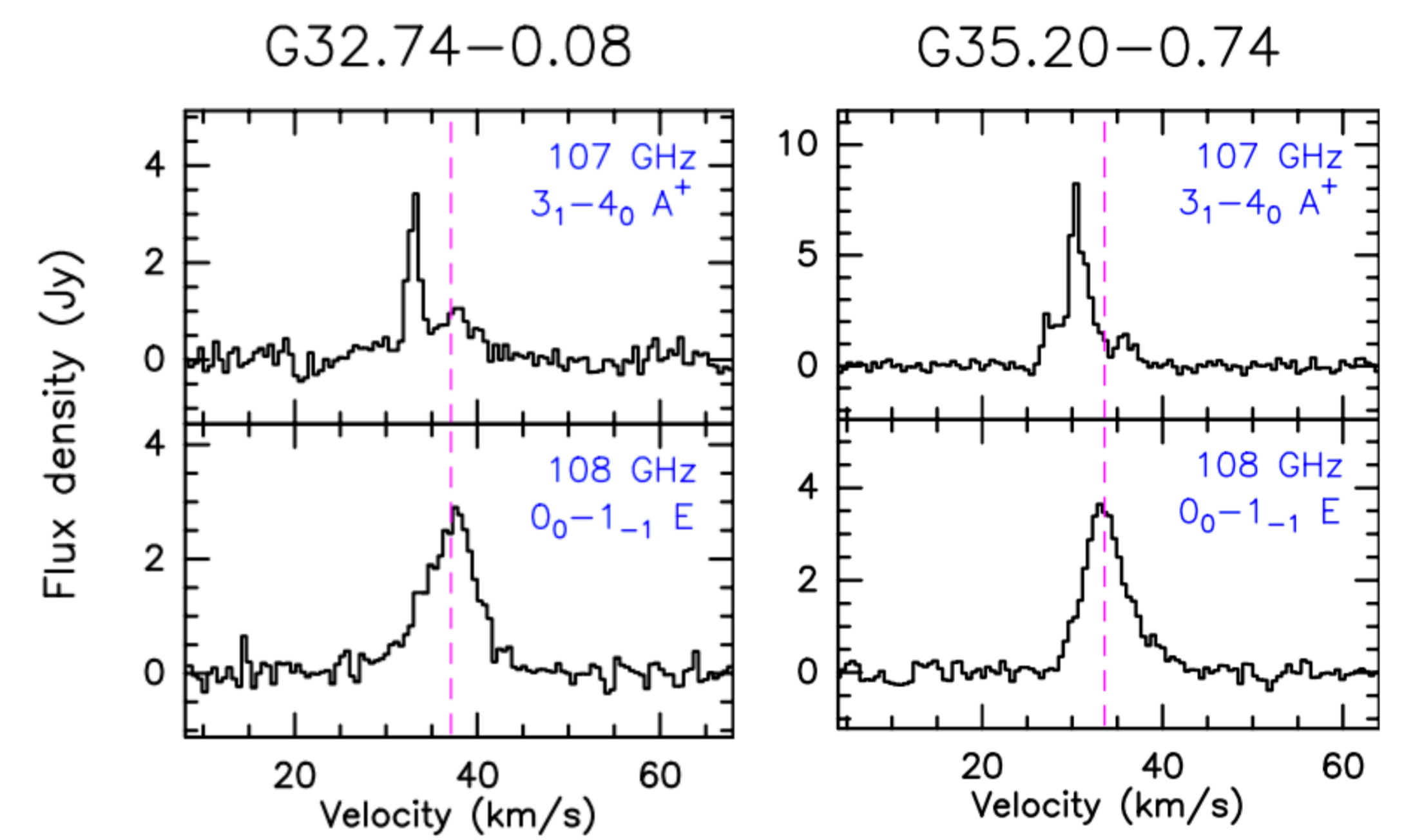


Fig. 2: Examples of sources showing newly detected 107 GHz methanol masers.

Anti-inversion of the 107 GHz line

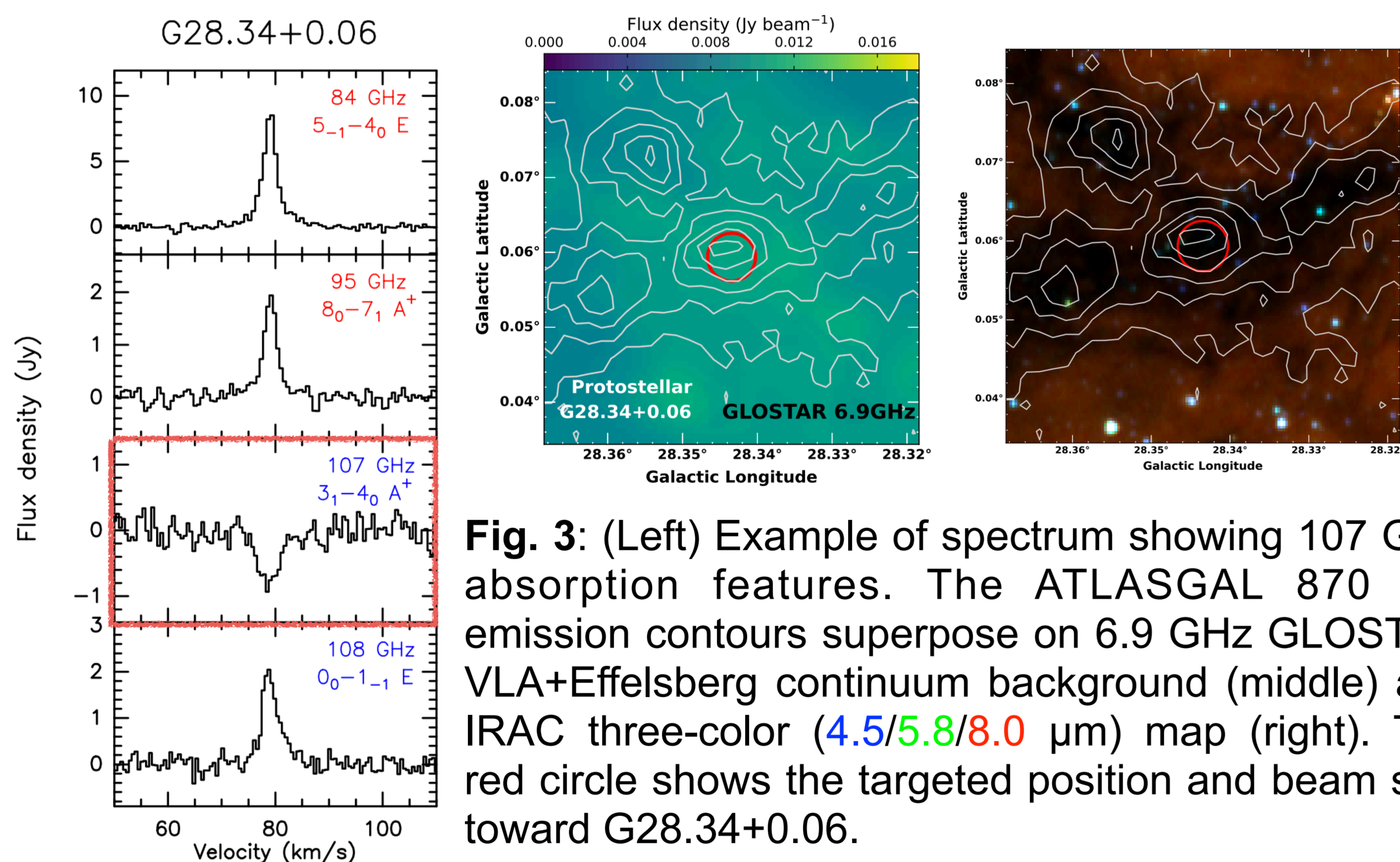


Fig. 3: (Left) Example of spectrum showing 107 GHz absorption features. The ATLASGAL 870 μ m emission contours superpose on 6.9 GHz GLOSTAR VLA+Effelsberg continuum background (middle) and IRAC three-color (4.5/5.8/8.0 μ m) map (right). The red circle shows the targeted position and beam size toward G28.34+0.06.

- **19** sources show 107 GHz CH₃OH absorption features
- Analysis on background dust and free-free emission \rightarrow 107 GHz CH₃OH absorption toward CMB, $T_{\text{ex}} < T_{\text{CMB}}$

The properties of class I CH₃OH masers are regulated by SiO traced shocks

- More and stronger class I masers were detected towards sources showing SiO line wings than sources without SiO wings.
- The total integrated intensity of class I masers is positively correlated with SiO integrated intensity and FWZP of SiO (2–1) emission.

Physical properties on a clump scale

- Isotropic luminosity or integrated intensity of class I masers vs. properties of associated ATLASGAL clumps:
positive correlation: L_{bol} , M_{clump} , $N(\text{H}_2)$
no statistically correlation: L/M , T_{dust} , $n(\text{H}_2)$
- Clumps with 104.3 GHz masers generally show brighter L_{bol} , warmer T_{dust} , larger L/M ratios, and denser environments.

The evolutionary stage of CH₃OH masers

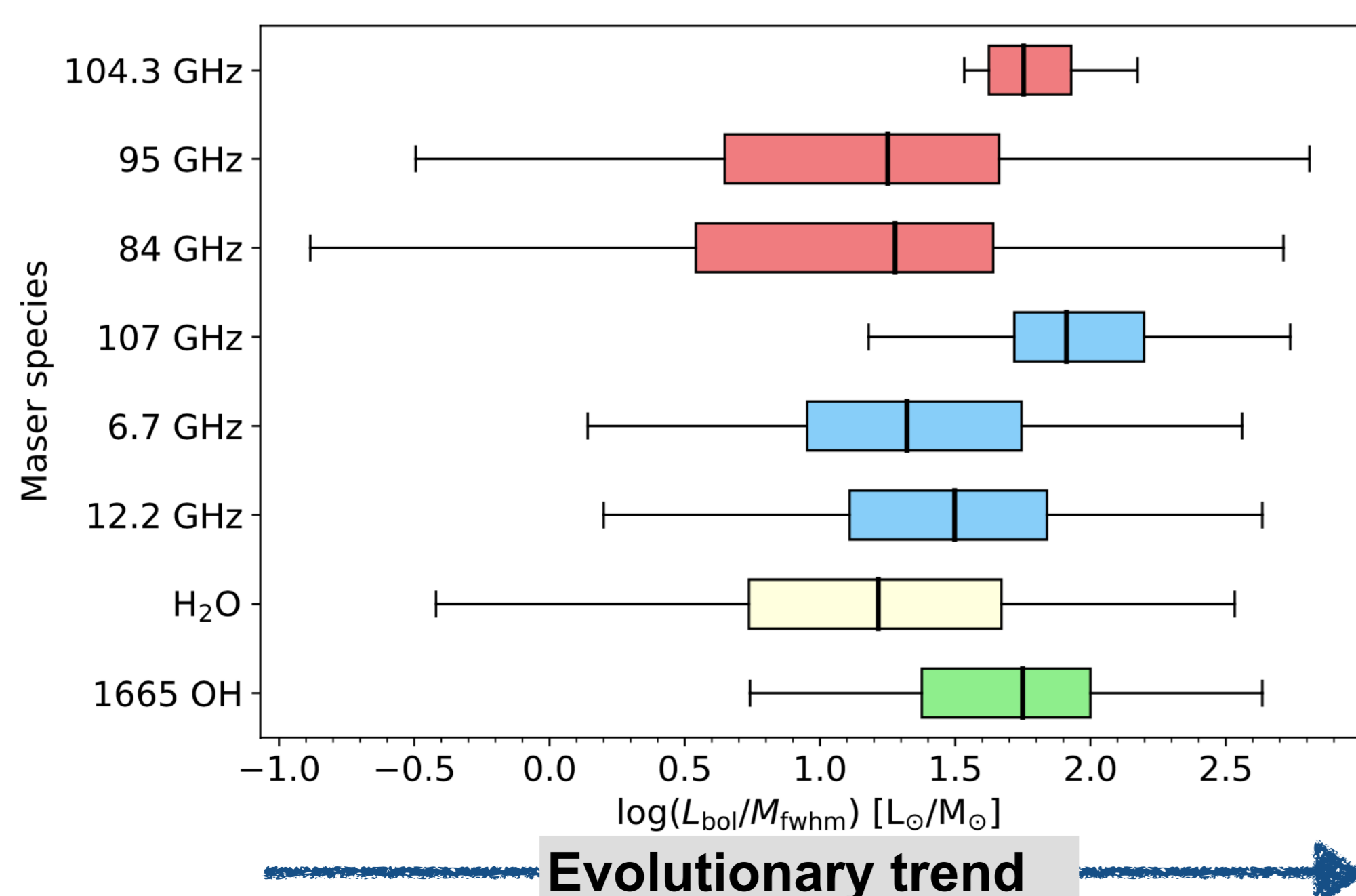


Fig. 4: Box plot shows the distributions of bolometric luminosity-to-mass ratios for clumps associated with different masers. Except for 3-mm CH₃OH masers, the plotting data are taken from Ladeyschikov et al. (2022).

Using multiple class I masers to constrain physical conditions

- myRadex (a RADEX analog)
- Maser feature at 15 km/s in G10.34-0.14 \rightarrow
 $T_{\text{kin}} = 57 \pm 3$ K
 $n(\text{para-H}_2) = 7.9(\pm 2.5) \times 10^5$ cm⁻³

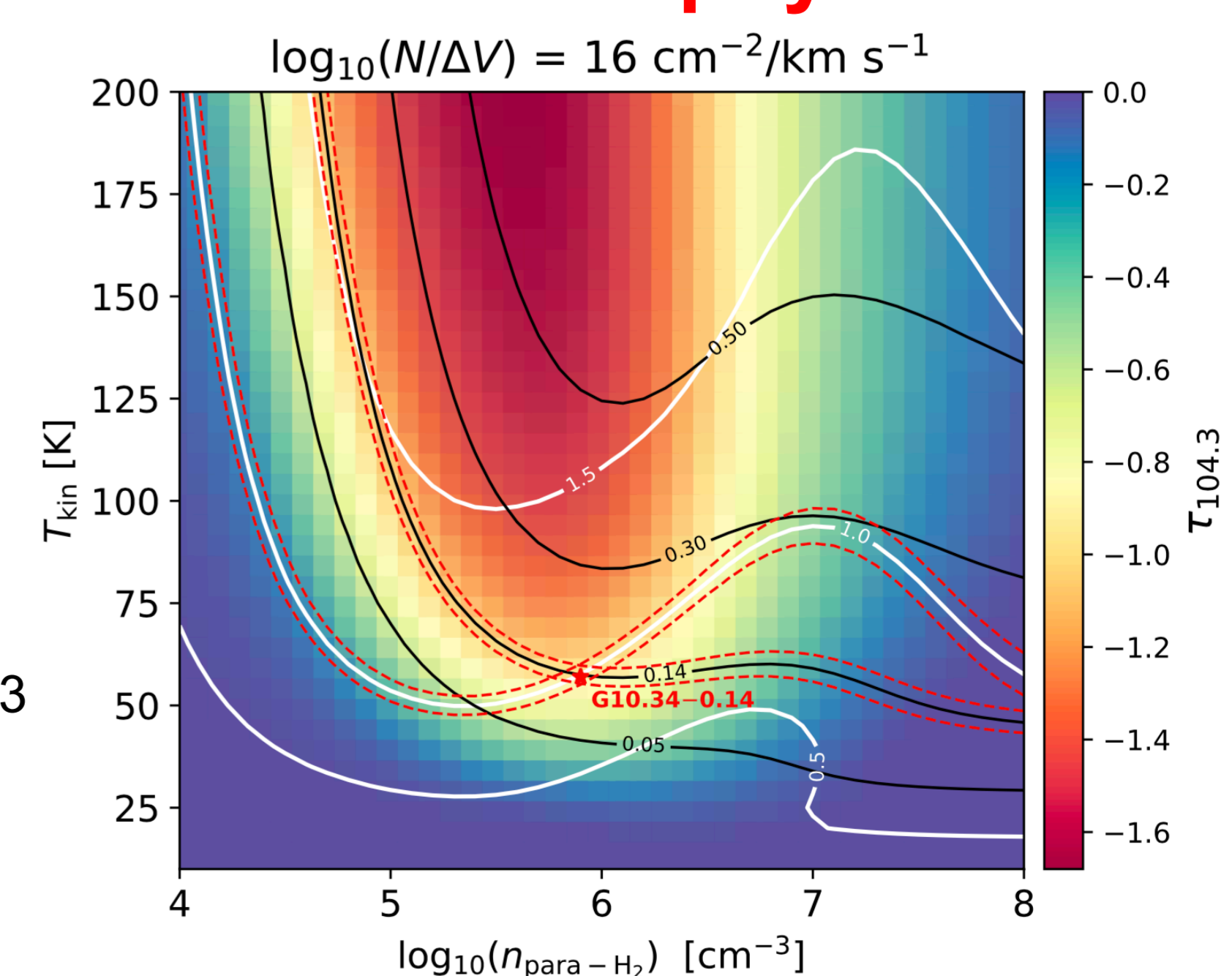


Fig. 5: Statistical equilibrium calculations result. The colored background shows negative optical depths of 104.3 GHz CH₃OH line. The black and white contours represent the peak intensity ratios of $T_{\text{r},104.3}/T_{\text{r},84}$ and $T_{\text{r},95}/T_{\text{r},84}$, respectively.

* W. Yang, et al. submitted to A&A; † W. Yang, et al. in prep.
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